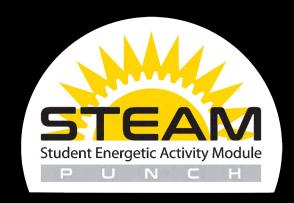
#### Polarimeter to Unify the Corona and Heliosphere

**Student Thermal Energetic Activity Module STEAM** 

> xc Galarraga Samantha Honan



**STEAM Science** July 7th, 2023



Amir Ca

Veronica Co

Craig







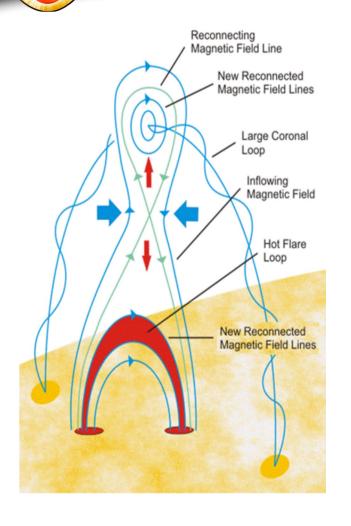
#### **Science Objectives**

Explore the enhancement of low First Ionization Potential (FIP) elements in the solar corona.

Explore how solar coronal plasmas are heated in flares and quiescent active regions.

Support PUNCH science in understanding the source regions of solar wind and coronal mass ejections.

## Magnetic Reconnection & Plasma Heating



Oppositely oriented field lines cancel

Field lines rearrange themselves into a lower energy state

Releases an explosion of energy

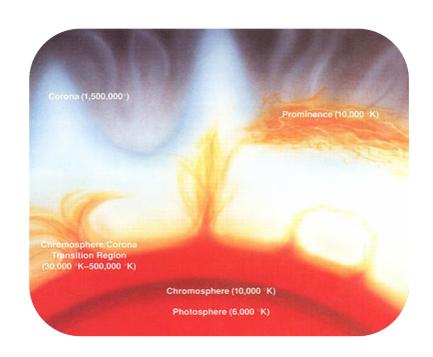
Releases heat and energy into the corona



#### Low-FIP Elements in the Corona

# Low FIP (< 10 eV) elemental abundances point to origin of plasma

- Prominent above thermal continuum
- Abundances enhanced by a factor of ~4 in corona over chromospheric values
- Abundances allow STEAM to infer origin of plasma for flares and active regions (AR)





# Why X-rays

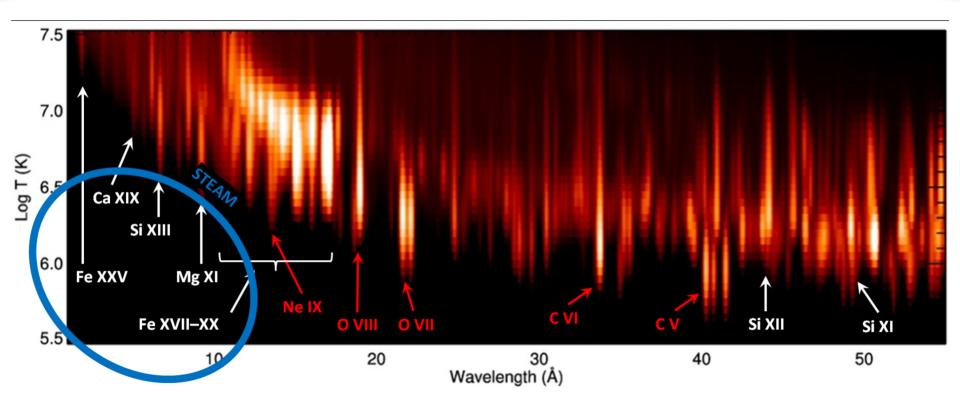
Soft X-ray Emissions Provide most direct signatures for hot plasmas Highly sensitive to hot temps & fluctuations

Not sensitive to cool temps

Hard X-ray Emissions Provides a more complete look at plasma evolution Covers higher energy events >10 keV

Extends thermal continuum Comprehensive coverage of thermal plasma temperature and composition





STEAM will be able to observe low FIP spectral lines with a greater resolution



# What is STEAM?

Slide updated

Soft X-ray Measurements



Silicon Drift Detector

Hard X-ray Measurements



Cadmium Telluride Detector

Expected Performance

SXR

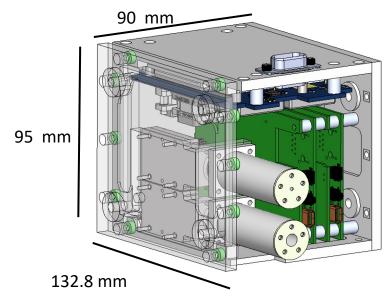
HXR

Performance		
Energy Range	1 to 7 keV	7 to 20 keV
Resolution	< 0.3 keV	< 1 keV
Field of View	5.25° to 10°	5.25° to 10°
Aperture Size	300 µm	2700 μm
Filter	PI, 5.5 μm	Al, 50 μm

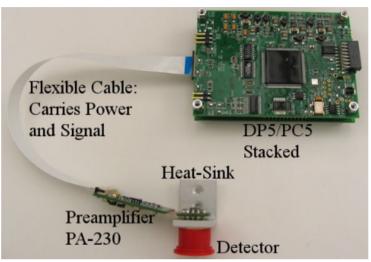




#### Current Model & Hardware



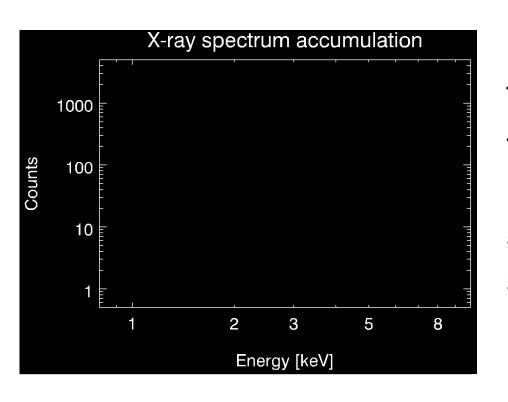
A rendering of the STEAM instrument and its dimensions



Detector head and electronics are about the size of a pack of cards.



## **Integrating Photon Counts Over Time**

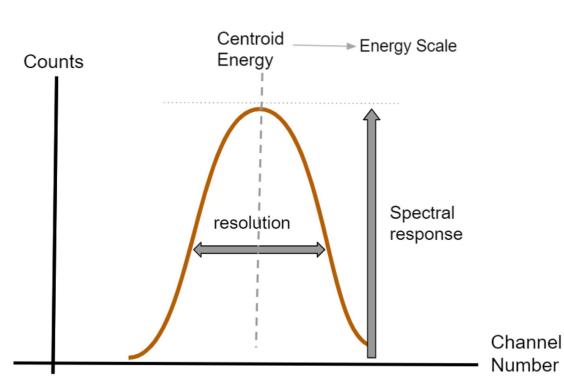


The individual histograms from 10-second integrations (red) can be summed to create a spectrum (white).



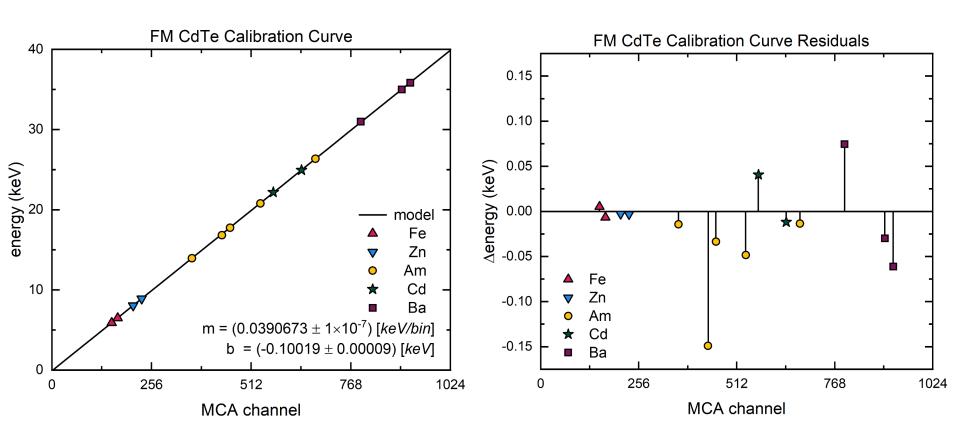
## Spectrometer Calibration

Below is a schematic of an emission line we would use to calibrate the energy scale of our detectors



On-ground calibration with X-ray emissions of radioactive isotopes Spectral Energy **Energy Scale** Resolution Response Attenuation Gain & Width of line of X-rays through Offset filters Absorption in Detector





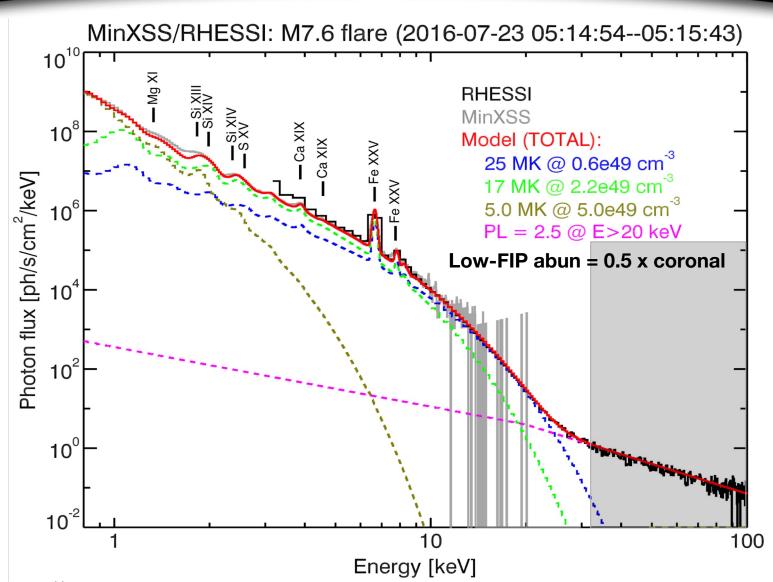


# Achieving Science Goals

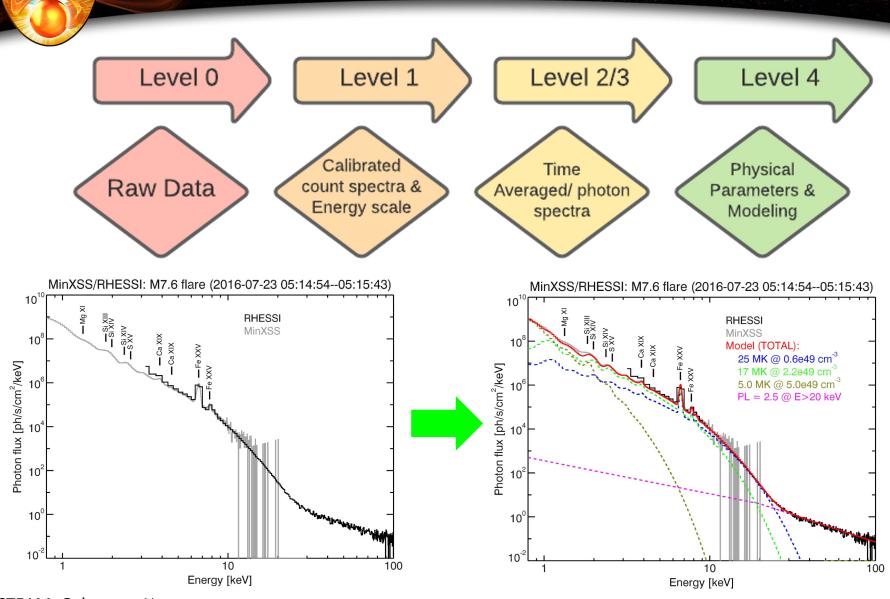
Observables	Applying the Physics	Modeling
	Define continuum and spectral line emissions	Continuum shape & line intensities
V roy photons from		•
X-ray photons from source		Forward Modeling
	Fit temperatures and abundances	
		Use Bremsstrahlung
	SXR (bound-bound	and atomic emission databases
▼	radiation)	databases
Respective energies	Majority of line emissions	
	HXR (free-free and free-bound radiation) Helps to constrain continuum shape	Chi-squared minimization to derive physical parameters



## Expected Data & Analysis/Modeling



# Tentative Data Pipeline





### Connections to PUNCH

CONNECTION BETWEEN
CORONA AND INNER HELIOSPHERE

ENERGY RELEASE
PROCESSES INTO CORONA

ORIGIN OF HEATED
PLASMA

TEMPERATURE OF LOW FIP ELEMENTS

SOLAR FLARES
& ACTIVE
REGIONS

**PUNCH** 

**STEAM** 



#### Current Work and Future Outlook

#### **Current Work:**

Finish assembly of Flight Model (FM)

FM Environmental tests

Analysis of FM spectrometer data for calibration between tests

#### **Future Outlook:**

Finish FM environmental tests

Complete final comprehensive testing

**Delivery in August!**